Cosmological Relaxation from Dark Fermion Production

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w/ Kenji Kadota, Ui Min, and Minho Son [1909.07706]

International joint workshop on the Standard Model and beyond

Introduction

- After Higgs was discovered, particle contents seem to be complete.
- Unsolved: naturalness, dark matter (DM), etc

$$m_h \ll M_P$$

Traditional sol: + sym, e.g. SUSY,
Higgs compositeness

@min: light new particles to be observed @ LHC?

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cosmological relaxation of EW scale

[Graham, Kaplan, Rajendran (GKR), 1504.07551]

Due to the null result @ LHC, this is an increasingly motivated scenario

Attempts to resolve downsides of GKR: e.g. **particle production (PP)**

Efficient way of dissipating energy

Various applications: e.g. preheating in reheating

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Efficient way of dissipating energy

Attempts to resolve **downsides of GKR**: e.g. **particle production (PP)**

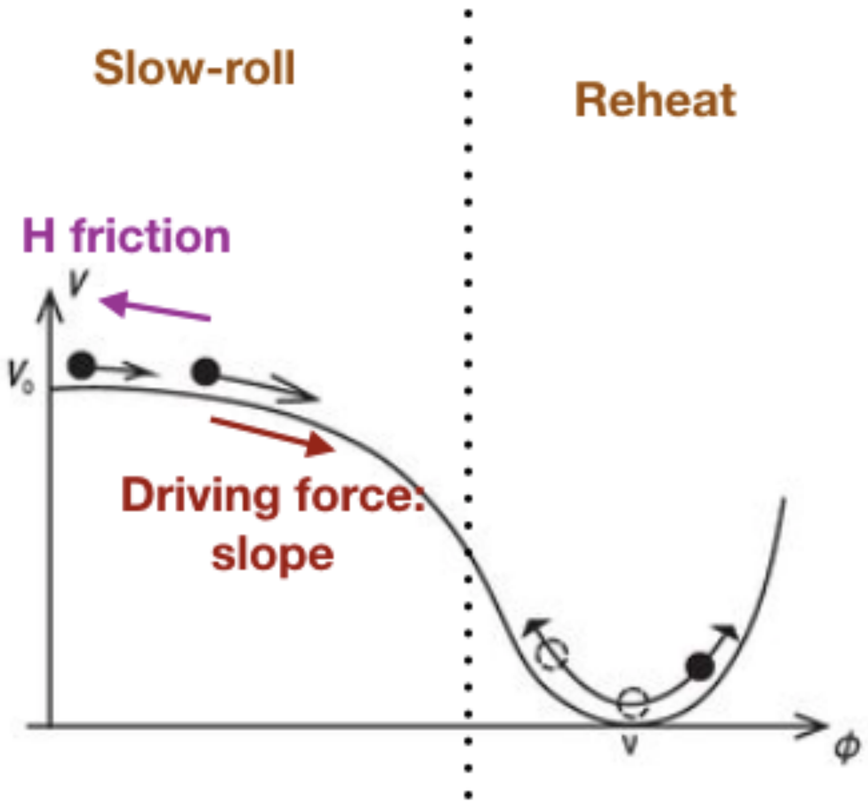
Various applications: e.g. preheating in reheating **Bose enhancement**

Pauli blocking usu. applications in fermion production (FP) quite limited

- **However** A recent application: **FP can support slow-roll inflation**

[Adshead, Pearce, Peloso, Robers, Sorbo, 1803.04501]

Traditional slow-roll inflation



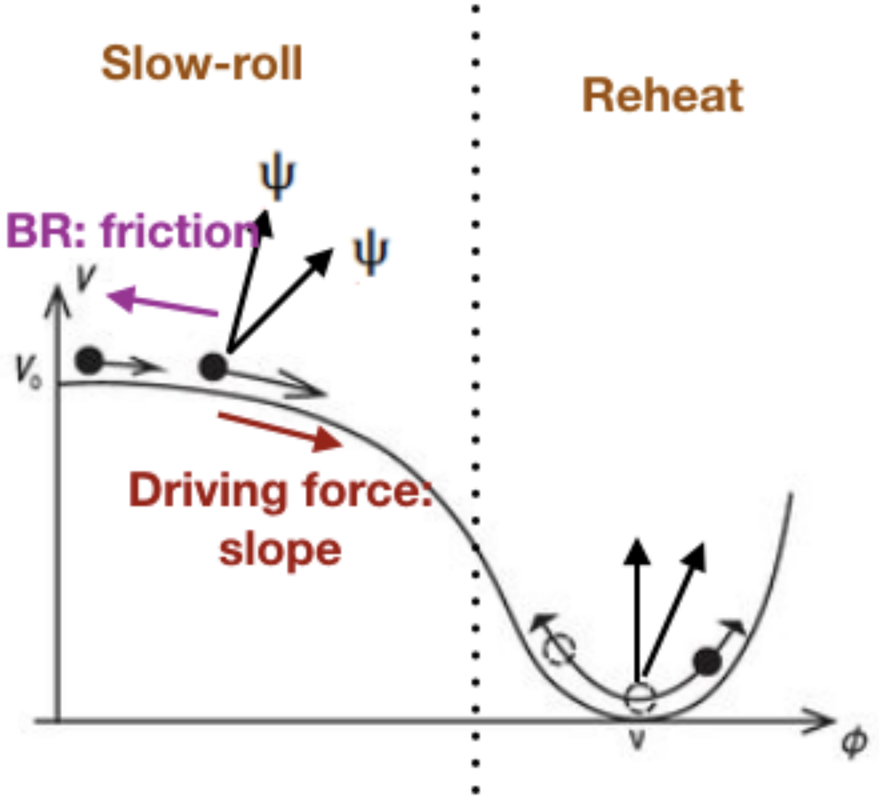
slope = Hubble friction

Potential has to be flat

$$\underset{\substack{\downarrow \\ \sim 0}}{\ddot{\phi}} + \boxed{3H\dot{\phi} + V_\phi(\phi)} = 0$$

FP supported slow-roll inflation

[Adshead, Pearce, Peloso, Robers, Sorbo, 1803.04501]



Backreaction (BR) = slope >> Hubble friction

Can support steep potential

$$\underset{\substack{\downarrow \\ \sim 0}}{\ddot{\phi}} + 3H\dot{\phi} + \boxed{V_\phi(\phi) = \mathcal{B}}$$

Introduction

Goals

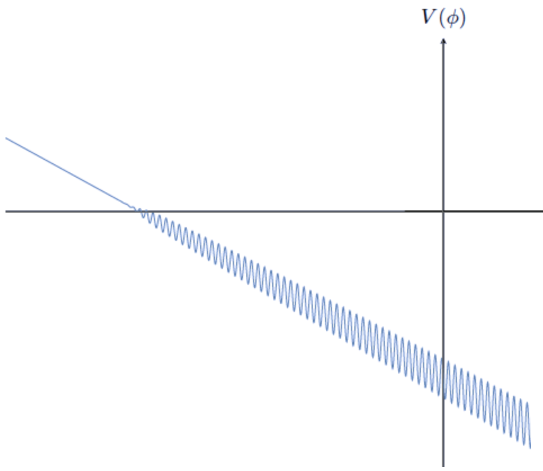
- Whether **FP** can be a **dominant way of dissipating relaxion energy** while maintaining the cutoff in a similar size to that of other variant models
- **Downsides of GKR gone**
- Phenomenology of this scenario: esp. if the **fermion is BSM**
-

GKR relaxation

- Relaxion: axion-like particle (ALP) whose periodic symmetry is **softly** and explicitly broken by a **small** coupling to Higgs (and also small self-coupling)

[Graham, Kaplan, Rajendran (GKR), 1504.07551]

- Smallness of Higgs mass: cosmological evolution



EW unbroken -> EW broken

$$\Delta V = \left(-\Lambda^2 + g\phi\right)|h|^2 + \left(g\Lambda^2\phi + \cdots\right) + \Lambda_c^4 \cos\left(\phi/f\right)$$

cutoff

small mass-dim coupling

Propto to Higgs vev <h>

Trapping: $g\Lambda^2 = \dfrac{\Lambda_c^4}{f}$

GKR relaxation

Conditions

$$\boxed{\Delta V = \left(-\Lambda^2 + g\phi\right)|h|^2 + \left(g\Lambda^2\phi + \cdots\right) + \Lambda_c^4 \cos\left(\phi/f\right)}$$

- Long enough slow-roll s.t. relation **can scan O(1) of its entire field range** => lower bound on number of efolds

Entire scan $\Delta\phi \sim \dot\phi\Delta t \sim \dot\phi\, N_e/H \underset{\uparrow}{\sim} (g\Lambda^2/H^2)\, N_e \gtrsim \Lambda^2/g \xrightarrow{\qquad\qquad} N_e \gtrsim H^2/g^2$

Slow-roll $3H\dot\phi + \frac{d\Delta V}{d\phi} \sim 0.$

- **Vac energy > change in the relation potential energy**

$$H^2 M_P^2 \gtrsim \Lambda^4$$

- **Barriers w/in Hubble sphere** $H^{-1} > \Lambda_c^{-1}$

slow-roll **Hubble time**

- **Classical > Quantum** $\Delta\phi \sim \dot\phi\Delta t \overset{\downarrow}{\sim} \frac{V'}{H}\,\frac{1}{H}\!\swarrow > H$

P.S. concerns regarding reheating

GKR relaxation

Problems $$\Delta V = \left(-\Lambda^2 + g\phi\right)|h|^2 + \left(g\Lambda^2\phi + \cdots\right) + \Lambda_c^4 \cos\left(\phi/f\right)$$

- QCD relaxion: O(f) shift of the local min of the QCD part => **O(1) theta parameter!** **Sol: + add. mech.**
- **Tiny coupling**: e.g. g~10^-31 GeV for QCD relaxion
- => exponentially large e-folds, super-Planckian excursion $\Delta\phi \geq \Lambda^2/g^2$

  $N_e \gtrsim H^2/g^2$

  Inefficient energy dissipation (Hubble)!

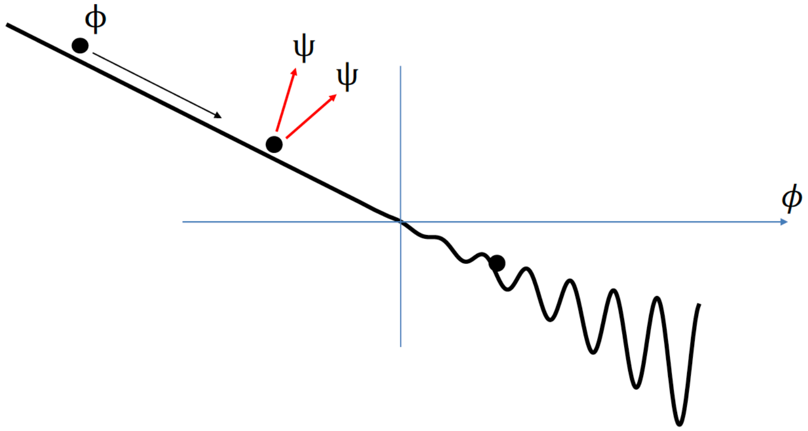
Sol: more efficient way of dissipating energy, e.g. particle production (PP) sourced by rolling relaxion

e.g. rolling relaxion exponentially producing gauge bosons to lose energy

- requiring a specific, nontrivial UV [Hook, Marques-Tavares (HMT), 1607.01786]

FP-supported relaxation?

Models



Fermion production

- Assume a flat FRW background

$$ds^2 = dt^2 - a^2 d\mathbf{x}^2 = a^2\left(d\tau^2 - d\mathbf{x}^2\right)$$

- Couple relaxion to fermion via derivative coupling

$$\Delta S = \int d^4x\sqrt{-g}\left[\bar{\psi}\left(i\,e^{\mu}{}_{a}\,\gamma^{a}D_{\mu} - m_{\psi} - \frac{1}{f_{\psi}}e^{\mu}{}_{a}\,\gamma^{a}\gamma^{5}\partial_{\mu}\phi\right)\psi\right]$$

Massless fermion => free field => production should be off

↓

If scanning starts in EW-sym phase,
the produced fermion can't be any SM fermion which is massless then.

The produced fermion must be BSM if scanning starts in EW-sym phase

Number operator **not well-defined** in this basis due to the **derivative coupling**

↓

[Adshead, Pearce, Peloso, Robers, Sorbo, 1803.04501]

[Min, Son, Suh, 1808.00939]

New basis $\psi \to a^{-3/2}\psi \quad \psi \to e^{-i\gamma^5\phi/f_\psi}\psi$

$$\Delta\mathcal{L} = \bar{\psi}\left(i\gamma^{\mu}\partial_{\mu} - m_R + i\,m_I\gamma^5\right)\psi \qquad \mathcal{H} = \bar{\psi}\left(-i\gamma^{i}\partial_{i} + m_R - i\,m_I\gamma^5\right)\psi$$

$$m_R = m_\psi a\,\cos(2\phi/f_\psi) \text{ and } m_I = m_\psi a\,\sin(2\phi/f_\psi).$$

Models

Strong back reaction supported slow-roll

$$\ddot{\phi} + 3H\dot{\phi} + V_\phi(\phi) = \mathcal{B}$$

$$\dot{\phi} \,\equiv\, \partial\phi/\partial t,\;\; V_\phi \,\equiv\, \partial V/\partial\phi$$

$$\mathcal{B} = \frac{2m_\psi}{fa^3}\langle\bar{\psi}\big[\sin\left(2\phi/f\right) + i\,\gamma^5\cos\left(2\phi/f\right)\big]\psi\rangle$$

$$\mu \equiv m_\psi/H \ll \xi$$

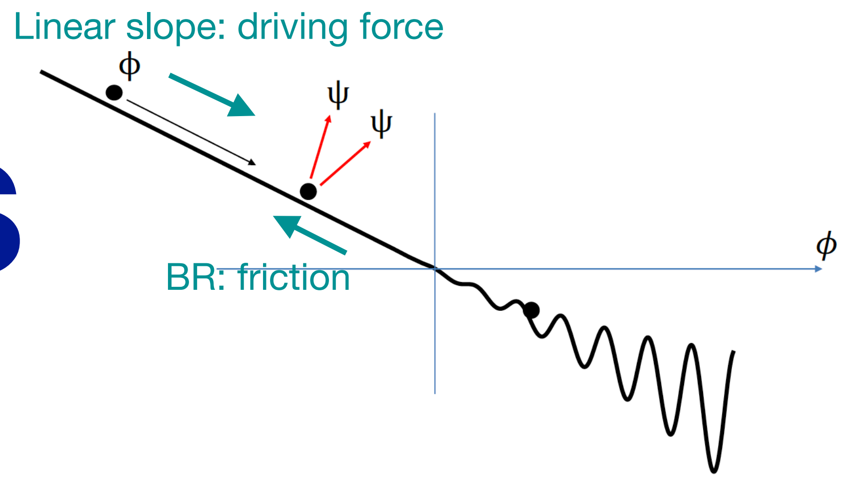
$$\mathcal{B} \sim -\frac{1}{f_\psi}H^4\mu^2\xi|\xi|$$

$$\boxed{\xi \equiv \frac{1}{2H}\frac{\dot{\phi}}{f_\psi}}$$

Fermions with heavier (but not too heavy) masses can also be produced,
but this simple expression for backreaction is no longer valid

Strong production: adiabaticity stongly violated
speed large enough, or coupling strength large enough

Models



Strong backreaction (BR) supported slow-roll

$$\ddot{\phi} + 3H\dot{\phi} + V_\phi(\phi) = \mathcal{B}$$

Strong back reaction

$$V_\phi(\phi)(= g\Lambda^2) \sim \mathcal{B} \quad \rightarrow \quad \dot{\phi} \sim 2\,\frac{g^{1/2}\Lambda f_\psi^{3/2}}{m_\psi} \sim \text{constant}$$

Sizable relaxion speed <=> not-too-small linear slope

Not-too-small g

$$\Delta V = \left(-\Lambda^2 + g\phi\right)|h|^2 + \left(g\Lambda^2\phi + \cdots\right) + \Lambda_c^4 \cos\left(\phi/f\right) + \frac{\partial_\mu \phi}{f_\psi} J_\psi^{5\mu}$$

Models

Constraints

- Slow-roll: Hubble friction must be small enough s.t. the slow-roll is maintained by slope compensated by the back reaction

$$V_\phi(\phi) > 3H\dot{\phi} \quad \rightarrow \quad m_\psi > 6\,\frac{H}{\Lambda}\frac{f_\psi^{3/2}}{g^{1/2}}$$

- Validity of EFT

$$\dot{\phi} \lesssim \Lambda^2 \quad \rightarrow \quad m_\psi \gtrsim 2\,\frac{g^{1/2}f_\psi^{3/2}}{\Lambda}$$

- Hidden fermion energy density small enough

$$\rho_\psi \sim 16\pi^2 H^4 \mu^2 \xi^3 \lesssim H^2 M_P^2 \rightarrow m_\psi \gtrsim \frac{\Lambda^3}{H^3}\frac{g^{3/2}f_\psi^{3/2}}{M_P^2}$$

Models

Constraints

- Relaxion kinetic energy < total energy

$$\dot{\phi}^2 \lesssim H^2 M_p^2 \quad \rightarrow \quad m_\psi \gtrsim 2\,\frac{\Lambda}{H}\frac{g^{1/2} f_\psi^{3/2}}{M_p}$$

Automatically when $\Lambda^4 \lesssim H^2 M_p^2$.

- Sufficient scanning, not-too-large efolding, sub-Planckian

$$\Delta\phi \gtrsim \frac{\Lambda^2}{g} \quad \rightarrow \quad m_\psi \lesssim 2N_e\,\frac{g^{3/2} f_\psi^{3/2}}{H\Lambda}$$

$$\Delta\phi = \dot{\phi}\Delta t = \dot{\phi}\,(N_e/H)$$

$$N_e \lesssim \mathcal{O}(10^{1\sim 3})$$

$$M_p > \Delta\phi \quad \rightarrow \quad m_\psi > 2N_e\frac{\Lambda}{H}\frac{g^{1/2} f_\psi^{3/2}}{M_p}$$

Models

Constraints

- Classical rolling > quantum spreading

$$\dot{\phi}\Delta t \gtrsim H \quad \rightarrow \quad m_\psi \lesssim \frac{g^{1/2} f_\psi^{3/2} \Lambda}{H^2}$$

- Barriers within Hubble sphere $H \lesssim \Lambda_c.$
- Precision of scanning $\Delta m_h^2 \sim g\Delta\phi \sim g\, 2\pi f \lesssim m_h^2$
- Temperature in the SM sector during scanning << v s.t. we're not scanning the thermal Higgs mass (ensured during inflation) (and we don't consider any fermion in a plasma to be produced by relaxion during scanning)

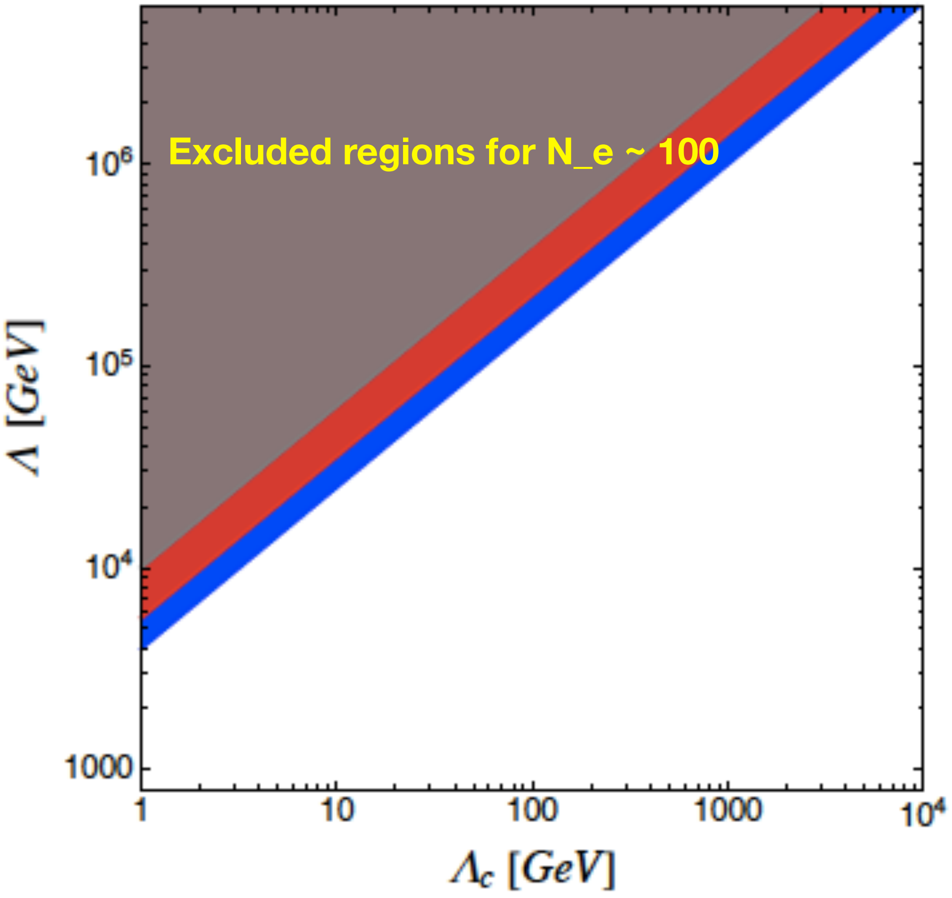
Models

Arrange all constraints in terms of bounds on fermion mass

Combine upper bounds and lower bounds on fermion mass

$$f > \Lambda \text{ and } H > \Lambda^2/M_p;$$

$$\Lambda < \min\left[(N_e/3)^{1/10} M_p^{1/5}\Lambda_c^{4/5} \ , \ (1/6)^{1/7} M_p^{3/7}\Lambda_c^{4/7} \ , \ N_e^{1/5}M_p^{1/5}\Lambda_c^{4/5} \right]$$



Focus on $\Lambda \sim 10^{4\sim5}$ GeV

Little hierarchy

Models

Forming periodic potential: model-dependent

Single extra scalar (relaxion)

- QCD relaxion: need extra mech. to solve strong CP, barrier height fixed
- **non-QCD relaxion**

2 extra scalars: double scanner

A single non-QCD relaxion

$$(\phi/f)\, G'_{\mu\nu} G'^{\mu\nu}$$

$$\Delta\mathcal{L}_{non\text{-}QCD} = m_L L L^c + m_N N N^c + y h L N^c + \tilde{y} h^\dagger L^c N$$

New fermions >~ EW scale

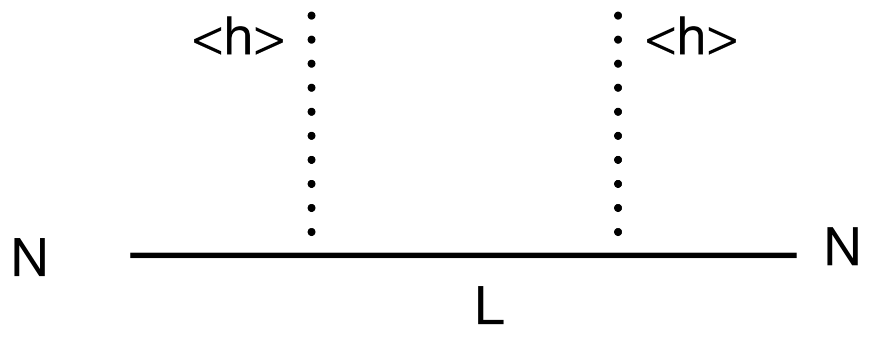
$$m_L \gg f_{\pi'} \gg m_N$$

Lighter fermion N responsible for forming condensate b/l confinement scale

$$m_N e^{i\phi/f} N N^c + \text{h.c.} = m_N N N^c \cos\frac{\phi}{f}$$

$$\langle N N^c \rangle \sim 4\pi f_{\pi'}^3$$

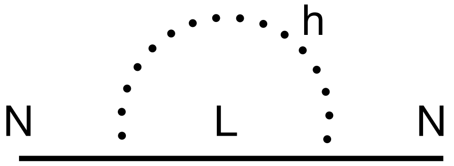
$$\Lambda_c^4 = 4\pi f_{\pi'}^3 m_N \sim 4\pi f_{\pi'}^3 \frac{y\tilde{y}\langle h\rangle^2}{m_L}$$



A single non-QCD relaxion

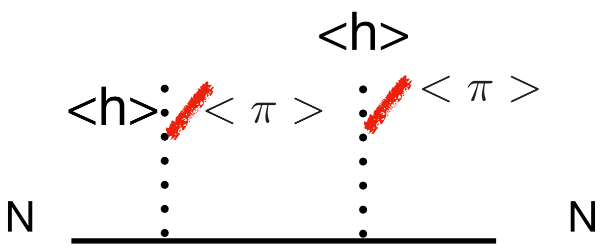
- For relaxation to work, **h-independent** contribution to the N mass must be **subdominant**

$$f_{\pi'} < \langle h \rangle \quad \text{and} \quad m_L < \frac{4\pi\langle h \rangle}{\sqrt{\log \Lambda / m_L}}$$



m_L ~ a few O(100) GeV

Constrained



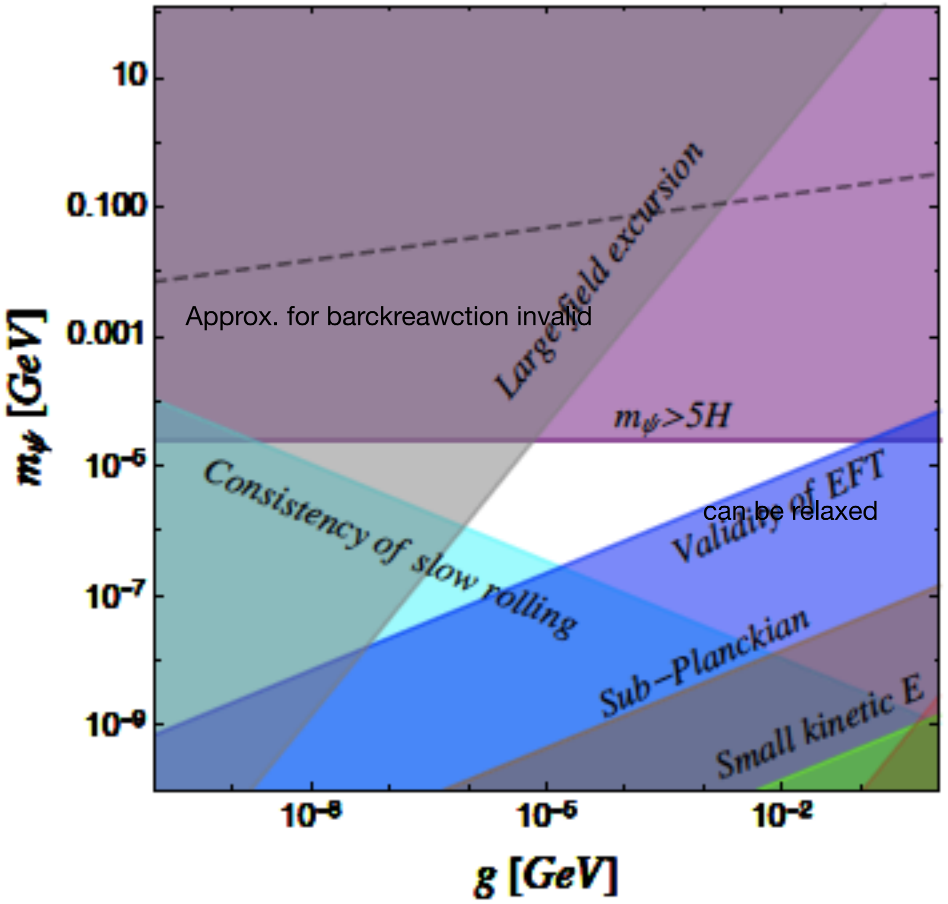
- N should be light enough compared to confinement scale

$$4\pi f_{\pi'} > \frac{y\tilde{y}\langle h \rangle^2}{m_L}$$

- EFT consistency: $f \gtrsim \Lambda$

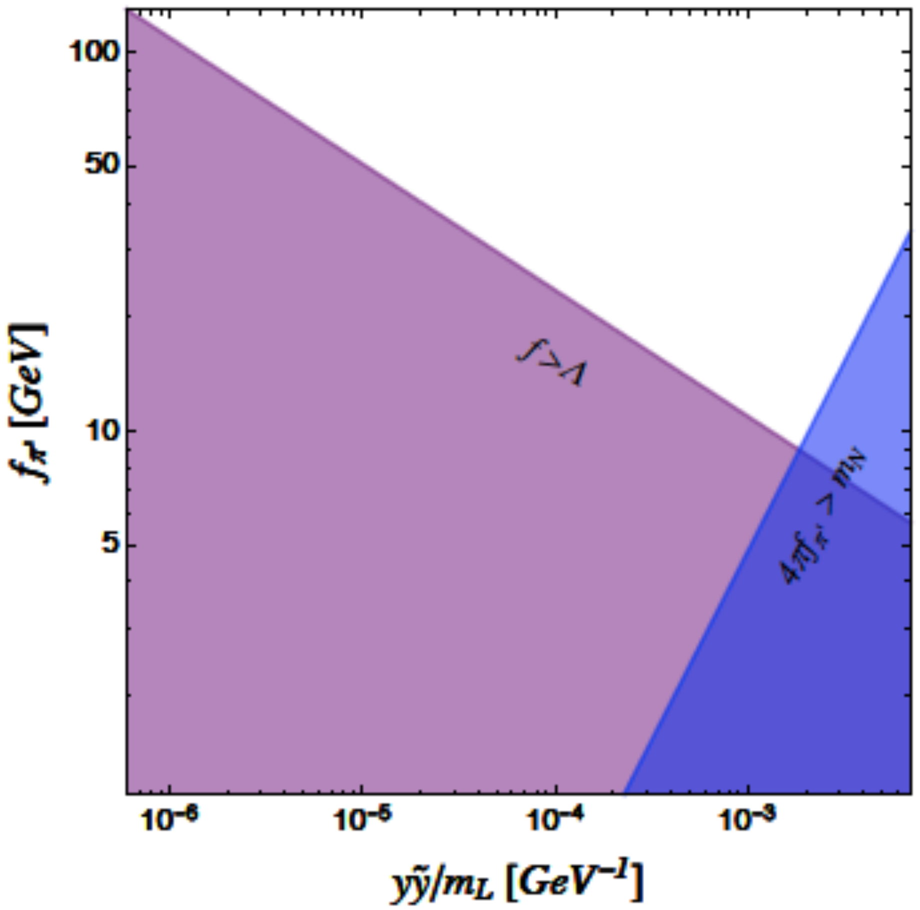
Other constraints:
Higgs decay, EWPT etc

A single non-QCD relaxion



Approx. for barckreawction invalid

can be relaxed



Λ	H	m_ψ	f_ψ	g	m_L	$y\tilde{y}$	$f_{\pi'}$	f	m_ϕ
10^4	5×10^{-6}	$1. \times 10^{-6}$	0.5	$1. \times 10^{-6}$	300	1.5×10^{-2}	45	3.4×10^4	$5. \times 10^{-2}$

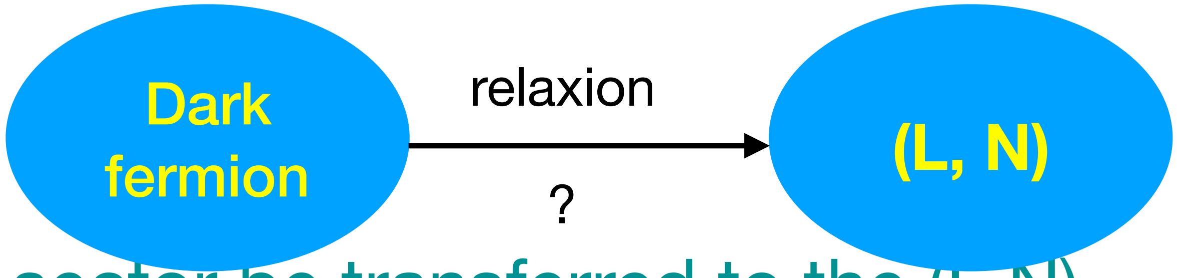
GeV
(except Yukawa)

$f_\psi \ll \Lambda$ Generic; to be solved separately, e.g. via clockwork

even worse in [Adshead, Pearce, Peloso, Robers, Sorbo, 1803.04501]

Assume: energy in the inflaton sector is smaller comparable to that of the relaxion-psi-SM-strong sector when the latter is in thermal eq.

A single non-QCD relaxion



- Can the energy in the fermion sector be transferred to the (L,N)-sector?

- Yes! Fermions thermalize relaxions. Thermal relaxions thermalize non-Abelian gauge bosons in (L,N)-sector.

:(

- Reheating temperature in (L,N)-sector may be high enough to erase the barriers! => 2nd rolling may ruin relaxation

:(

- Very non-trivial constraints to prevent this to happen

-

-> “Double scanner”

Double scanner

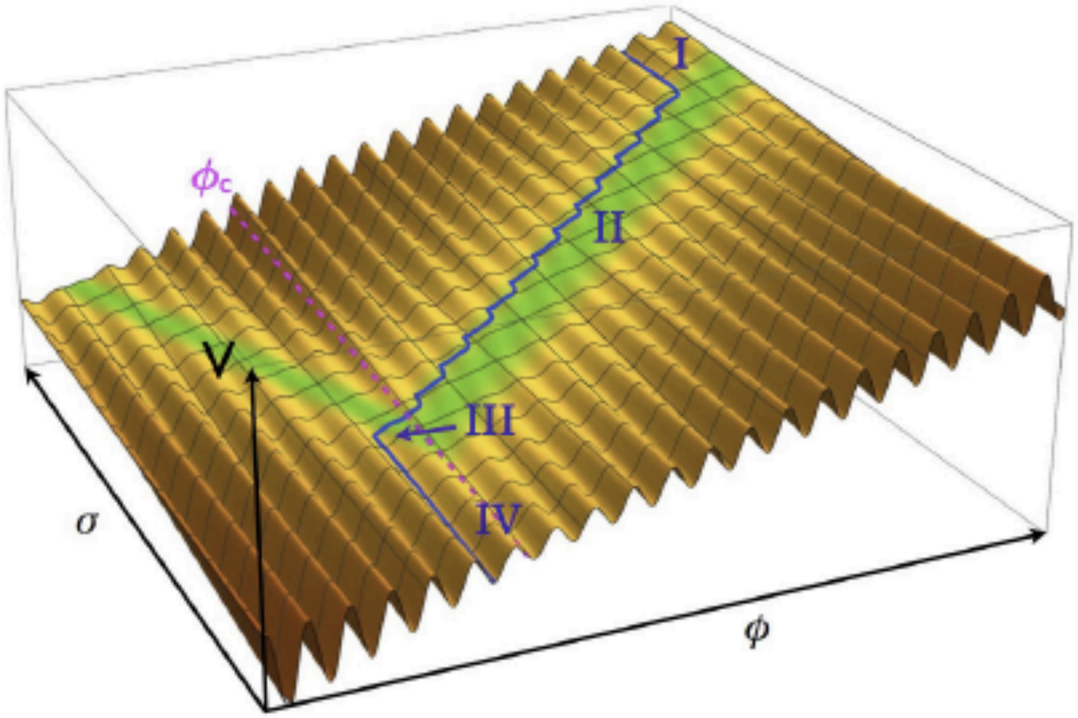
[Espinosa, Grojean, Panico, Pomarol, Pujolas, Servant, 1506.09217]

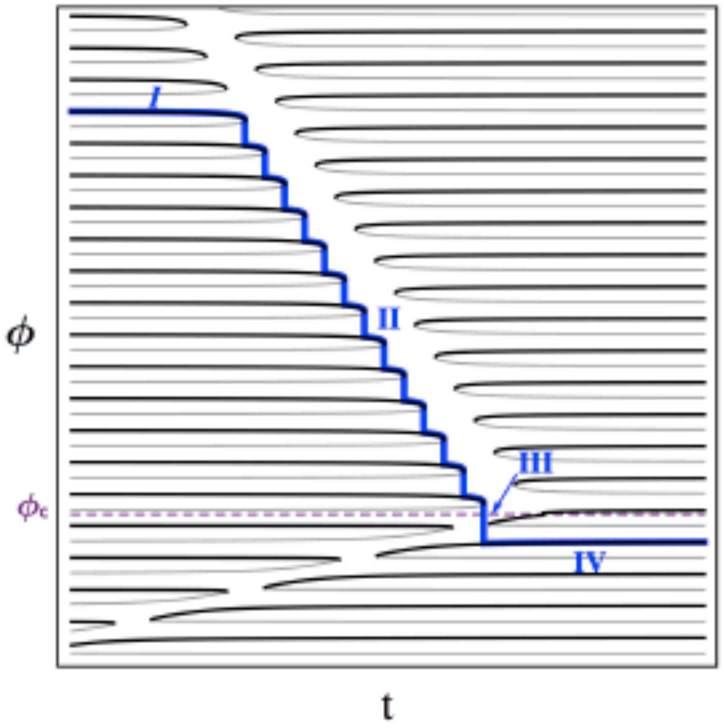
Confinement scale ~ cutoff

Barriers won’t be erased :)

$$\Delta V = g\Lambda^2\phi + g_\sigma\Lambda^2\sigma + \left(-\Lambda^2 + g\phi\right)|h|^2 + A(\phi,\,\sigma,\,h)\cos\left(\phi/f\right)$$
$$+\frac{\partial_\mu\phi}{f_\psi}\bar{\psi}\gamma^\mu\gamma^5\psi + \frac{\partial_\mu\sigma}{f_\sigma}\bar{\psi}\gamma^\mu\gamma^5\psi + m_\psi\bar{\psi}\psi\ ,$$
$$A(\phi,\,\sigma,\,h) = \epsilon\Lambda^4\left(\beta + c_\phi\frac{g\phi}{\Lambda^2} - c_\sigma\frac{g_\sigma\sigma}{\Lambda^2} + \frac{|h|^2}{\Lambda^2}\right)$$

[Figures from 1506.09217]

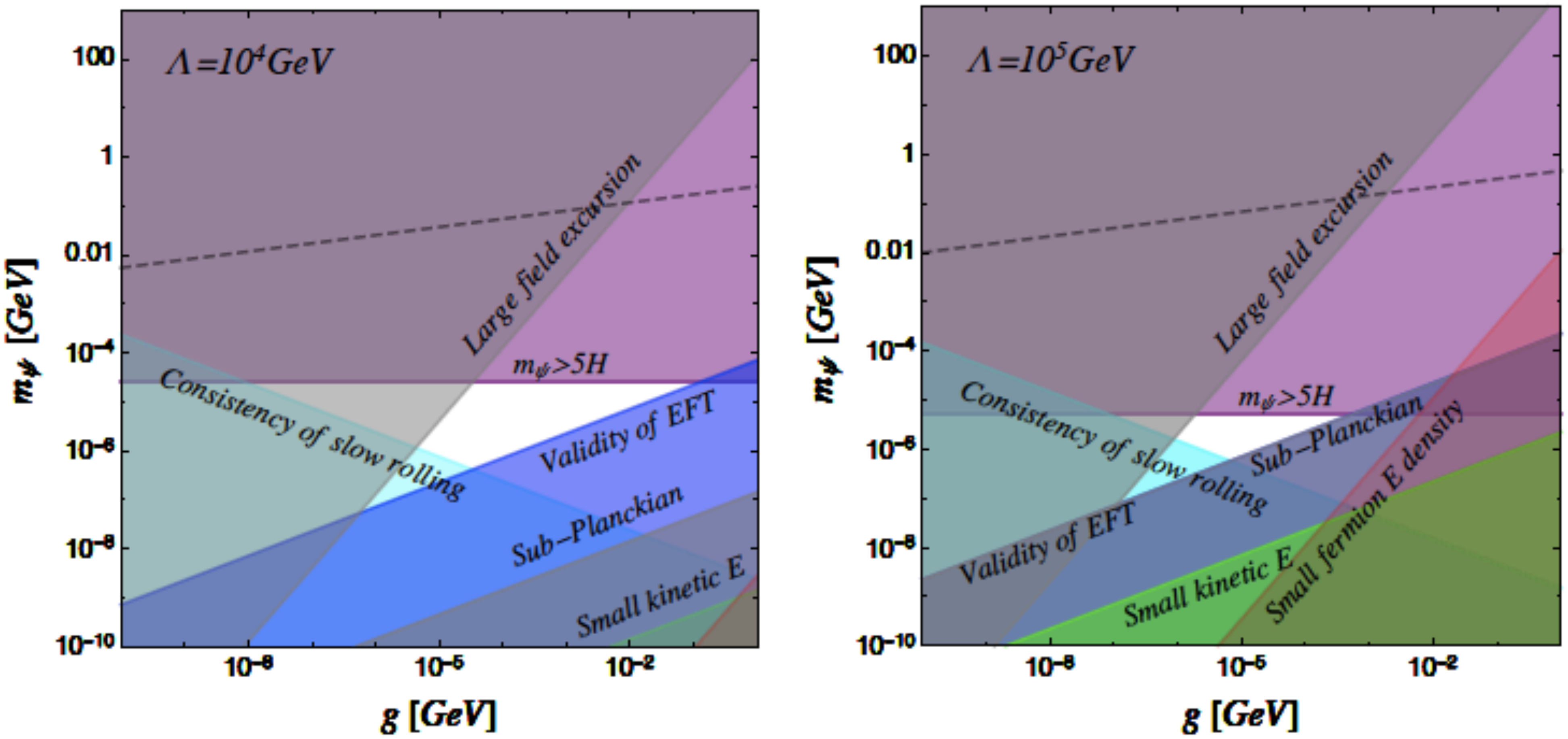




Conditions need to be satisfied for the mech. to work…

Double scanner

$$f_\psi = f_\sigma$$



Λ	H	m_ψ	$f_\psi \sim f_\sigma$	g	$g_\sigma(\sim m_\sigma)$	ϵ	f	m_ϕ
10^4	5×10^{-6}	$1. \times 10^{-6}$	0.5	$1. \times 10^{-5}$	$2. \times 10^{-6}$	$1. \times 10^{-5}$	6.1×10^4	5.2
10^5	1×10^{-6}	$1. \times 10^{-6}$	5	$1. \times 10^{-6}$	$2. \times 10^{-7}$	$2. \times 10^{-6}$	1.2×10^5	1.2×10^2

GeV
(except
epsilon)

Prospect for Dark Matter

Relic abundances for scalars

Benchmark pt $m_\phi \sim \mathcal{O}(100)$ GeV

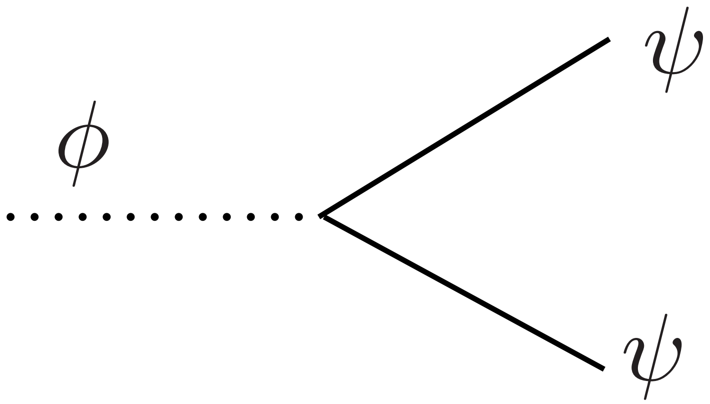
$m_\psi \sim$ KeV

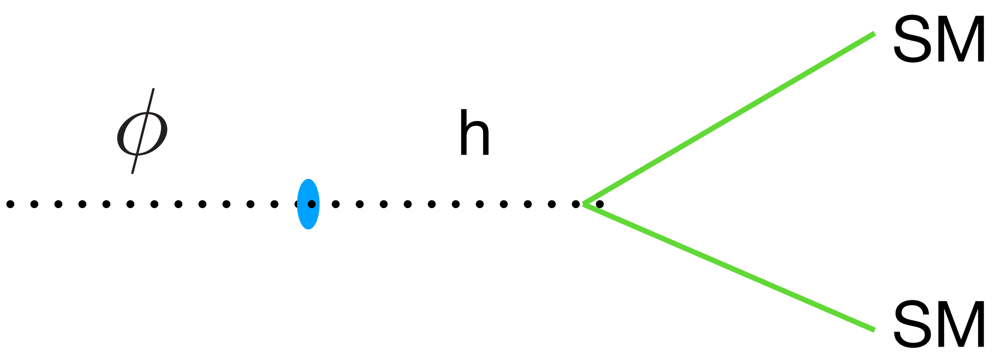
ϕ mixing w/ h, w/ mixing angle $\theta_{\phi h} \sim 2gv/m_h^2.$ can decay into hidden fermion before BBN

$$\Gamma_\phi = \theta_{\phi h}^2 \Gamma_h(m_\phi) + \Gamma_{\phi\to\psi\psi}(m_\phi)$$

$$m_\phi^2 = \frac{\epsilon\Lambda^4}{f^2} \sim \frac{g}{v^2}\frac{\Lambda^4}{f} = \frac{g}{f}\left(\frac{\Lambda}{v}\right)^4 v^2$$

$$\Gamma_{\phi\to\psi\psi} = \frac{1}{2m_\phi}\frac{8m_\psi^2 m_\phi^2}{f_\psi^2}\frac{1}{8\pi}\sqrt{1-\frac{4m_\psi^2}{m_\phi^2}} = \frac{1}{2\pi}\frac{m_\psi^2}{f_\psi^2}m_\phi\sqrt{1-\frac{4m_\psi^2}{m_\phi^2}}$$





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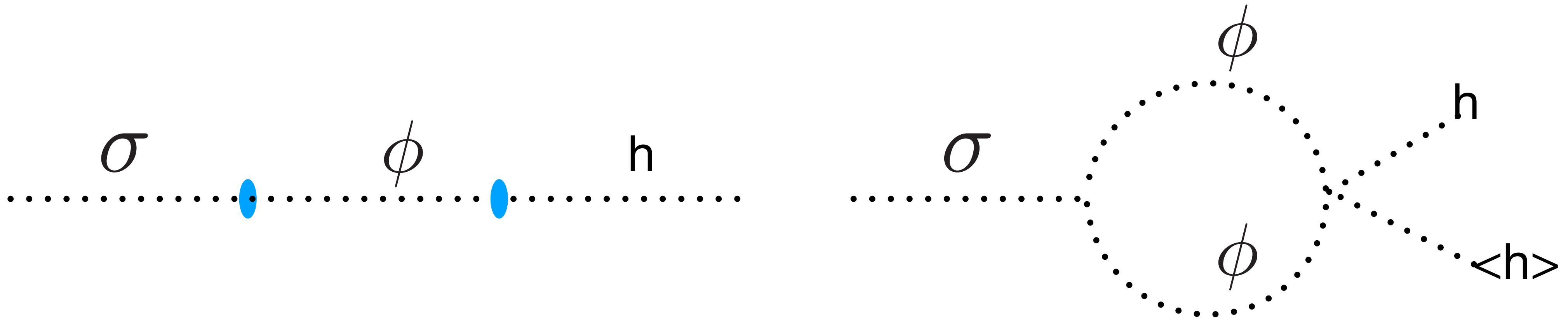

σ $\theta_{\sigma\phi} \sim \frac{g_\sigma f v^2}{\Lambda^4},\ \theta_{\sigma h} \sim \mathrm{Max}\left(\theta_{\sigma\phi}\theta_{\phi h},\ \frac{g^2}{16\pi^2}\frac{g_\sigma \Lambda^4}{f^2 v^3 m_h^2}\right)$

$m_\sigma^2 \sim g_\sigma^2$

$m_\sigma \sim\ \mathrm{KeV\ for\ our\ benchmark\ pt}$

If $\sigma \to \psi\psi$ is turned on, decay into hidden fermion before BBN

If $\sigma \to \psi\psi$ is turned off, only decay into SM fields vis h-mixing w/ small rate

Need to worry about
its abundance

Non-thermal:
misalignment

$m_\sigma^2(\Delta\sigma)^2 \quad \Delta\sigma \sim \sqrt{N_e}H$

$$\Omega_0^\sigma = \frac{\rho_0^\sigma}{\rho_c} \sim \frac{1}{\rho_c} m_\sigma^2 N_e H^2 \left(\frac{T_0}{\sqrt{m_\sigma M_p}}\right)^3 \ll 1 \qquad T_{osc} = \sqrt{m_\sigma M_p}$$

Prospect for Dark Matter

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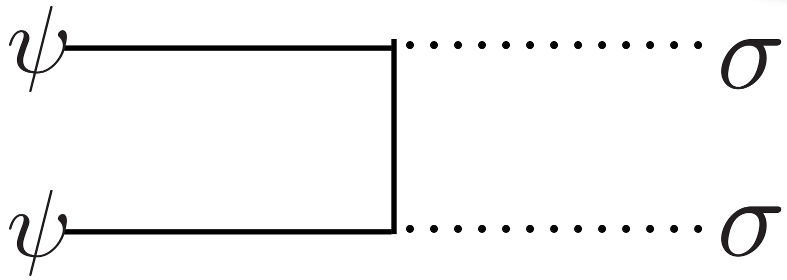
$m_\sigma^2 \sim g_\sigma^2$

$m_\sigma \sim$ KeV for our benchmark pt



σ $\theta_{\sigma\phi} \sim \frac{g_\sigma f v^2}{\Lambda^4}, \; \theta_{\sigma h} \sim \mathrm{Max}\left(\theta_{\sigma\phi}\theta_{\phi h}, \; \frac{g^2}{16\pi^2}\frac{g_\sigma \Lambda^4}{f^2 v^3 m_h^2}\right)$

If $\sigma \to \psi\psi$ is on, decay into hidden fermion before BBN

If $\sigma \to \psi\psi$ is off, only decay into SM fields vis h-mixing w/ small rate

Need to worry about its abundance

Thermal: $\Gamma_{\psi\psi\to\sigma\sigma}(T) \sim \frac{m_\psi^2}{f_\psi^4}T^3$ $T_d \sim \frac{f_\psi^4}{M_p}\frac{1}{m_\psi^2}$

More precise calculation:

$T_0 \sim \mathcal{O}(10^{-13})$ GeV

$$T_d \sim 10^{-4}\mathrm{GeV} \gg m_\sigma \xrightarrow{\text{decoupled while relativistic}} \Omega_0^\sigma \sim m_\sigma\, T_0^3\, g_{*S}(T_0)/\rho_C\, g_{*S}(T_d) \sim \mathrm{O}(10)$$

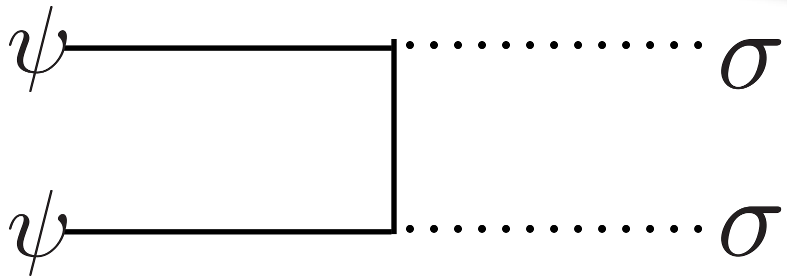
Over!

Prospect for Dark Matter

Relic abundances for scalars

$m_\sigma^2 \sim g_\sigma^2$

$m_\sigma \sim$ KeV for our benchmark pt



σ $\theta_{\sigma\phi} \sim \frac{g_\sigma f v^2}{\Lambda^4}, \ \theta_{\sigma h} \sim \mathrm{Max}\left(\theta_{\sigma\phi}\theta_{\phi h}, \ \frac{g^2}{16\pi^2}\frac{g_\sigma \Lambda^4}{f^2 v^3 m_h^2}\right)$

If $\sigma \to \psi\psi$ is on, decay into hidden fermion before BBN

If $\sigma \to \psi\psi$ is off, only decay into SM fields vis h-mixing w/ small rate

Need to worry about its abundance

Thermal: $\Gamma_{\psi\psi\to\sigma\sigma}(T) \sim \frac{m_\psi^2}{f_\psi^4}T^3$ $\quad T_d \sim \frac{f_\psi^4}{M_p}\frac{1}{m_\psi^2}$ decoupled while relativistic

Non-universal coupling:

$T_0 \sim \mathcal{O}(10^{-13})$ GeV

$f_\psi \to f_\sigma = \Lambda \gg f_\psi \qquad T_d \sim \mathcal{O}(10^{1-2})\mathrm{GeV}$

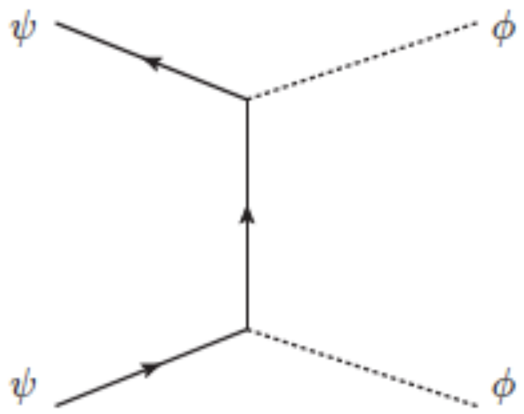
$$\Omega_0^\psi \sim \frac{m_\psi T_0^3}{\rho_c}\frac{g_{*S}(T_0)}{g_{*S}(T_d)} \times \frac{1}{S}$$

Entropy dilution

$\Omega_0^\sigma \sim m_\sigma T_0^3\, g_{*S}(T_0)/\rho_C\, g_{*S}(T_d)$ ~ O(1)

Prospect for Dark Matter

Relic abundances for the fermion



Benchmark pt:

$$m_\phi \sim \mathcal{O}(100)\ \mathrm{GeV}$$
$$m_\psi \sim\ \mathrm{KeV}$$

Hidden fermion decouples @ T ~ O(100) GeV
(not able to produce relaxion on shell => 2-step stops to be valid;
chain process not in thermal equilibrium)

Hidden fermion decouples while highly relativistic

$$\Omega_0^\psi \sim \frac{m_\psi T_0^3}{\rho_c}\frac{g_{*S}(T_0)}{g_{*S}(T_d)} \times \frac{1}{S}$$

~ O(0.1)

alpha-Lyman: $m_{WDM} \gtrsim 5\mathrm{keV} \times \left(\frac{g_{*S}(T_d) \sim \mathcal{O}(10^2)}{g_{*S}(T \ll \mathrm{MeV})}\right)^{-1/3} \sim 1\mathrm{keV}$

Thermal Warm Dark Matter (WDM)?

Summary

- Achieved cosmological **relaxation with back reaction from hidden fermion production**

- **Downsides of GKR all gone**

- The models require a **relatively strong coupling between the relaxion and the hidden fermion** => seemingly EFT inconsistency? Explained by clockwork etc?

- Possible thermalization b/t produced fermions and relaxion even during inflation. Double scanner: thermalized relaxion can’t thermalize the visible sector during inflation.

-

Thank you!

Backup

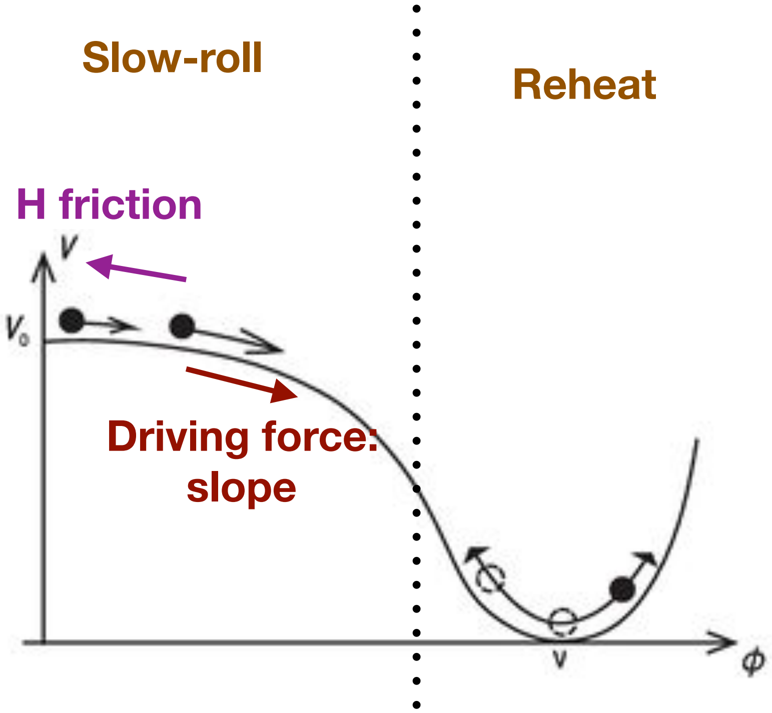
Introduction: particle production

- Particle production is an efficient way of dissipating energy

- Various applications in pheno and cosmology

- Exponentially producing bosons: example in reheating: preheating

Traditional slow-roll inflation



Slow-roll

Reheat

H friction

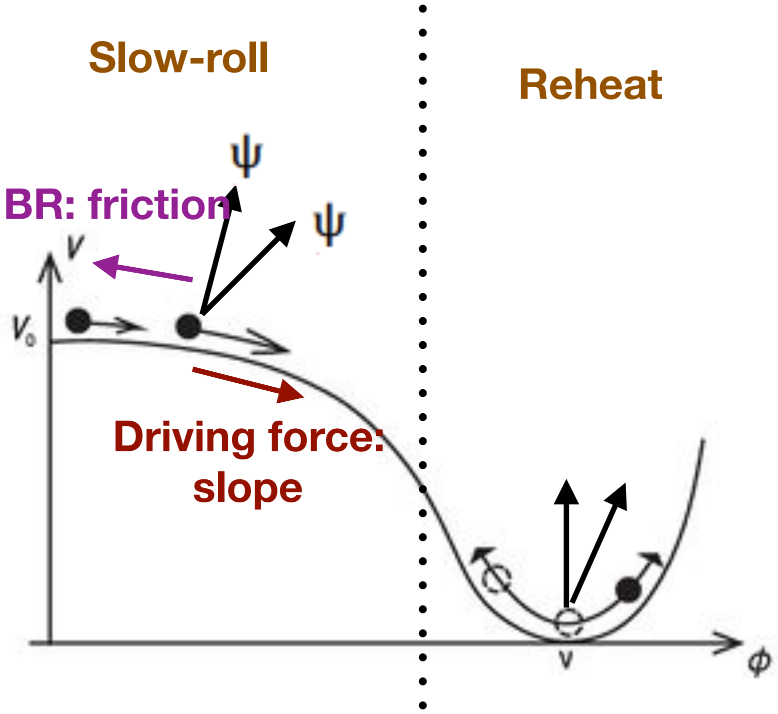
Driving force:
slope

slope = Hubble friction

Potential has to be flat

FP supported slow-roll inflation

[Adshead, Pearce, Peloso, Robers, Sorbo, 1803.04501]



Backreaction (BR) = slope >> Hubble friction

Can support steep potential

Introduction: relaxation

$$\Delta V = \left(-\Lambda^2 + g\phi\right)|h|^2 + \left(g\Lambda^2\phi + \cdots\right) + \Lambda_c^4 \cos\left(\phi/f\right)$$

cutoff

small mass-dim parameter

Propto to Higgs vev <h>

By NP effect, QCD or non-QCD

- Initial: relaxion has a very large field value (s.t. **positive Higgs mass-squared**) and **slowly rolls** down from its potential

$$\phi \gtrsim \Lambda^2/g \qquad \mu^2 \equiv -\Lambda^2 + g\phi > 0$$

- Rolling of relaxion => scanning Higgs mass
- At some pt: Higgs mass = 0. After this pt, <h> starts to develop, height of the periodic barrier increases
- When the height of the barrier is enough to compensate the linear slope and trap the relaxion, <h> is set to the correct EW VEV v.

Stopping condi: linear slope matches the barrier slope

Introduction: relaxation

Conditions

- After relaxion stops rolling and **reheating** occurs, the reheating temperature must be low enough s.t. barriers don't melt or the traveling distance of the **2nd rolling** leads to a change in Higgs mass smaller than EW scale

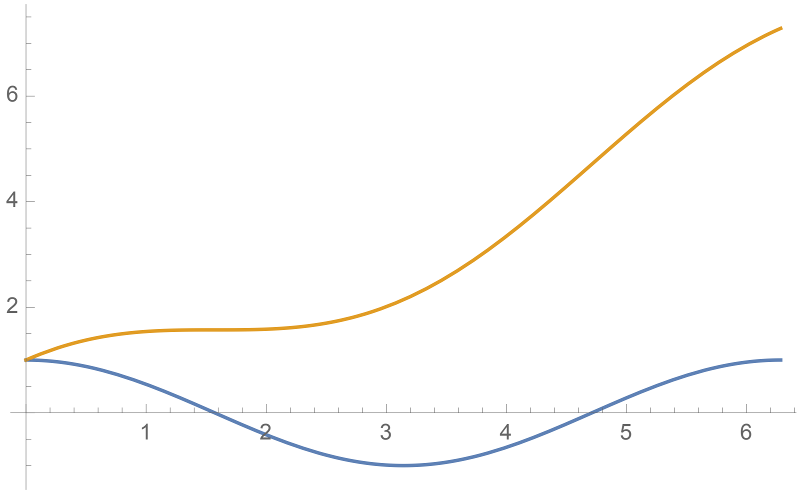
Introduction: relaxation

Problems

$$\Delta V = \left(-\Lambda^2 + g\phi\right)|h|^2 + \left(g\Lambda^2\phi + \cdots\right) + \Lambda_c^4\cos\left(\phi/f\right)$$

- QCD relaxion: O(f) shift of the local min of the QCD part => **O(1) theta parameter!** => Sol: + additional mech. (e.g. a separate inflaton)

$$\begin{aligned}&\text{Near } \phi \sim \Lambda^2/g,\\ &\quad \Delta V \sim g\Lambda^2\phi + \Lambda_c^4\cos(\phi/f)\\ &\quad\quad \sim \Lambda_c^4\left[\frac{\phi}{f} + \cos\left(\frac{\phi}{f}\right)\right]\end{aligned}$$

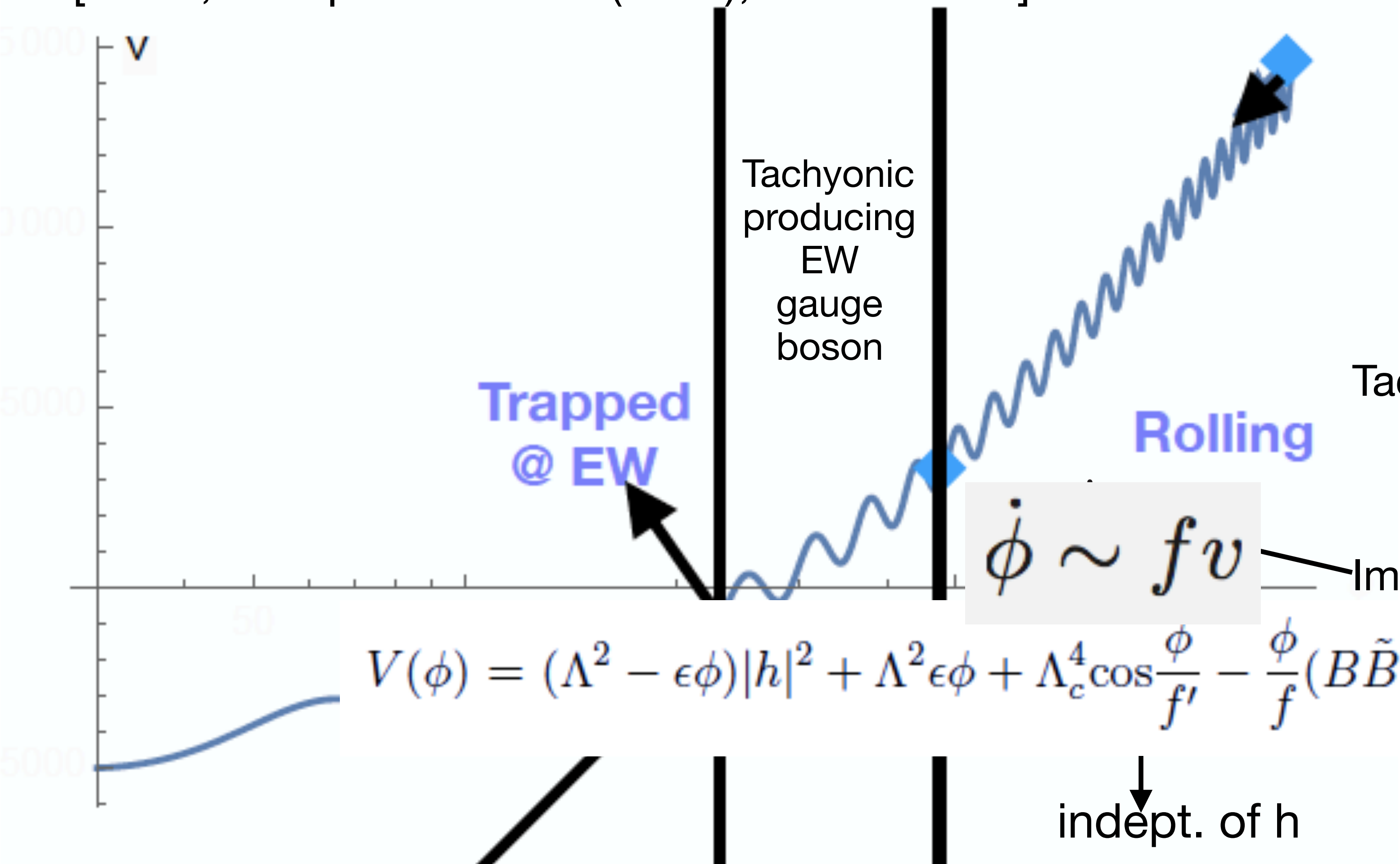


Introduction: particle production

- Exponentially producing bosons: example in relaxion models: tachyonic production of gauge bosons to stop relaxion

Fixed barrier height; h-dependence in the cond. to trigger tachyonic production

[Hook, Marques-Tavares (HMT), 1607.01786]



$\dot{\phi} \sim fv$

$$V(\phi) = (\Lambda^2 - \epsilon\phi)|h|^2 + \Lambda^2\epsilon\phi + \Lambda_c^4 \cos\frac{\phi}{f'} - \frac{\phi}{f}(B\tilde{B} - W\tilde{W}),$$

indept. of h

Start:
EW-broken
with ***very negative*** Higgs mass-square;
sufficient speed to pass min

Tachyonic production of gauge bosons when

$$\ddot{A}_\pm + (k^2 + m^2 \pm k\frac{\dot{\phi}}{f})A_\pm = 0 \quad \dot{\phi} \geq fm$$

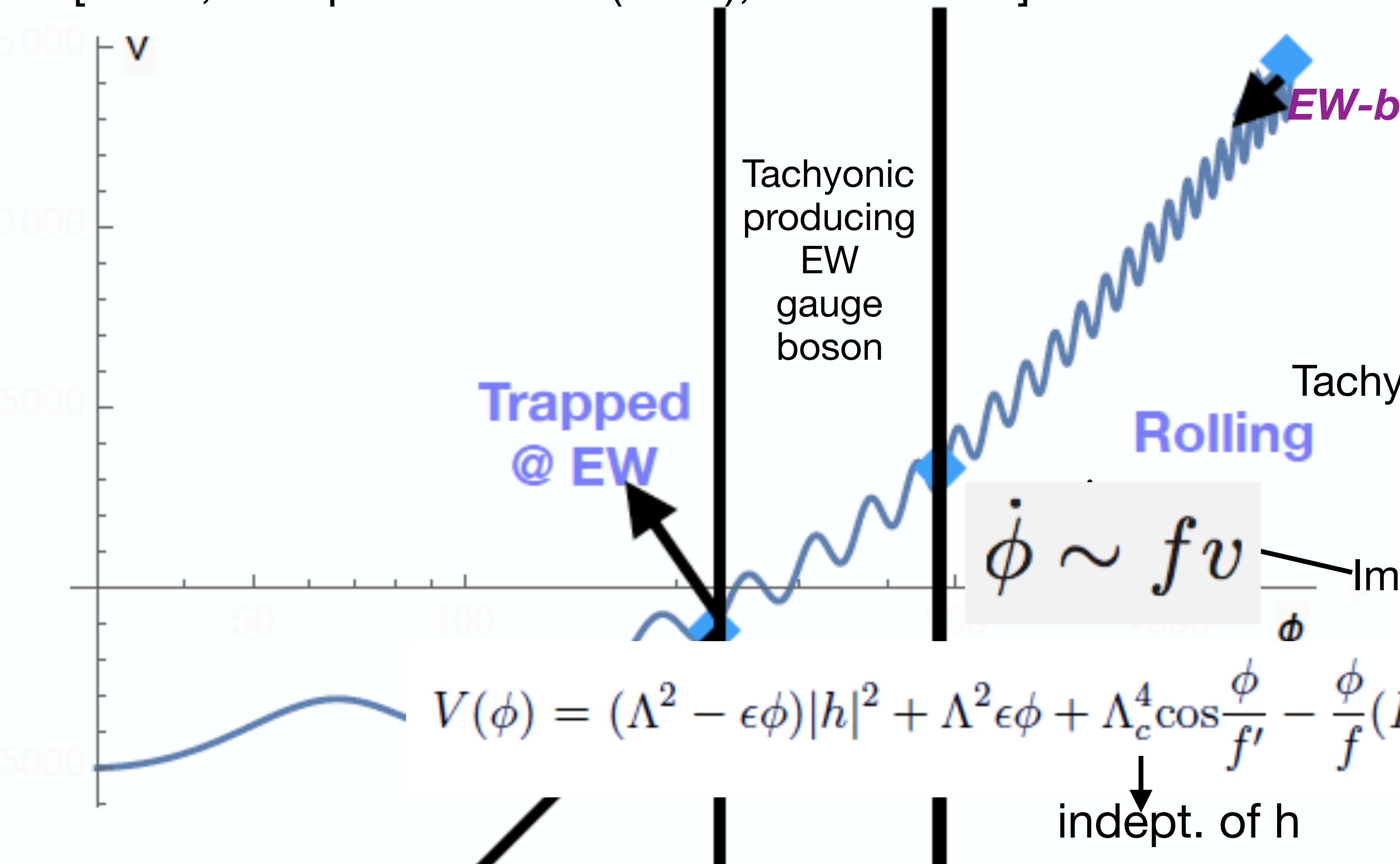
Impose: tachyonic production starts at **m ~ v**

Introduction: particle production

- Exponentially producing bosons: example in relaxion models: tachyonic production of gauge bosons to stop relaxion

Fixed barrier height; h-dependence in the cond. to trigger tachyonic production

[Hook, Marques-Tavares (HMT), 1607.01786]



$$\dot{\phi} \sim fv$$

$$V(\phi) = (\Lambda^2 - \epsilon\phi)|h|^2 + \Lambda^2\epsilon\phi + \Lambda_c^4\cos\frac{\phi}{f'} - \frac{\phi}{f}(B\tilde{B} - W\tilde{W}),$$

indept. of h

Start:
EW-broken phase with ***very negative*** Higgs mass-square
relatively fast relaxion ***speed*** to pass min

Scan <h> => scan gauge boson mass m

Tachyonic production of gauge bosons exists when

$$\ddot{A}_\pm + (k^2 + m^2 \pm k\frac{\dot{\phi}}{f})A_\pm = 0 \quad \dot{\phi} \geq fm$$

Impose: tachyonic production starts at **m ~ v**

Tachyonic production quickly
drains relaxion energy
and stabilizes it

Introduction: particle production

HMT

- HMT solved problems in GKR. A specific UV needed.
- Cutoff in HMT: <~ 10^{4~5} GeV
- Tachyonic production is so strong that the slow-roll can't really be maintained ("quasi-slow-roll")
- Can we **maintain slow-roll with particle production** as the friction?

Introduction: particle production

- **Fermion production**?
- **Can’t be exponential** due to Pauli blocking
- But may be **sufficient to support** a (“steep-slope”) **slow-roll**

[Adshead, Pearce, Peloso, Robers, Sorbo, 1803.04501]

Potential slope ~ fermion back reaction >> Hubble term

Introduction

Goals

- To maintain slow-roll
- To explain (little) hierarchy
- To use **fermion production** as the major friction
- No extremely small parameter
- To explore pheno. of such a scenario, eps. of the fermion produced by relaxion if it’s BSM

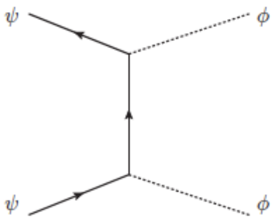
Inflaton: containing the most energy during inflation,
assumed to be a separate sector s.t. not too much
inflationary dynamics involved

Models

A single non-QCD relaxion

- One more problem: reheating temperature in the (L,N)-sector can't be too high

- Strong u/t-channel to make relaxion thermal



For $T \gg m_\phi, m_\psi,$

if such interactions are in eq.

$$\Gamma \sim \frac{m_\psi^2}{f_\psi^4} T^3 > H \sim \frac{T^2}{M_P}$$

- $\Rightarrow$

$$T > \frac{f_\psi^4}{M_P m_\psi^2} \sim 10^{-6}\text{GeV}\left(\frac{f_\psi}{1\text{GeV}}\right)^4 \left(\frac{10^{18}\text{GeV}}{M_P}\right)\left(\frac{10^{-6}\text{GeV}}{m_\psi}\right)^2$$

Assume:
inflaton energy dilutes faster than radiation after inflation
s.t. now the universe is radiation-dominated

Double-scanner

$$\Delta V = g\Lambda^2\phi + g_\sigma\Lambda^2\sigma + \left(-\Lambda^2 + g\phi\right)|h|^2 + A(\phi,\,\sigma,\,h)\cos\left(\phi/f\right)$$
$$+ \frac{\partial_\mu\phi}{f_\psi}\bar{\psi}\gamma^\mu\gamma^5\psi + \frac{\partial_\mu\sigma}{f_\psi}\bar{\psi}\gamma^\mu\gamma^5\psi + m_\psi\bar{\psi}\psi\ ,$$

$$A(\phi,\,\sigma,\,h) = \epsilon\Lambda^4\left(\beta + c_\phi\frac{g\phi}{\Lambda 2} - c_\sigma\frac{g_\sigma\sigma}{\Lambda 2} + \frac{|h|^2}{\Lambda 2}\right)$$

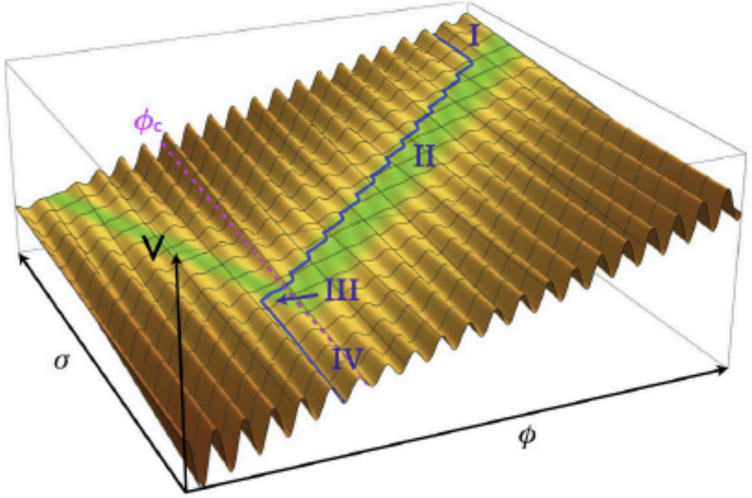
Constraints

I: sigma rolling, relaxion trapped $A \sim \epsilon\Lambda^4$

II: relaxion needs to scan before reaching the critical pt. $A \sim 0$

$$d\phi(t)/d\sigma(t) = (g/g_\sigma)^{1/2} > d\phi_*/d\sigma \qquad c_\phi\, g^{3/2} > c_\sigma\, g_\sigma^{3/2}$$

$$g > g_\sigma \text{ for } c_\phi \sim c_\sigma \sim \mathcal{O}(1).$$



III: Relaxion exits the trajectory (periodic slope << linear slope) to evolve along the path where A grows as h grows $A \sim \epsilon\Lambda^2 h^2$

$$d\phi(t)/d\sigma(t) < d\phi_*/d\sigma \qquad (c_\phi - 1/(2\lambda))\, g^{3/2} > c_\sigma\, g_\sigma^{3/2}$$

Relaxion trapped when slope condi.

$$g\Lambda^2 = \frac{A}{f} \sim \frac{\epsilon\Lambda^2 v^2}{f}.$$

IV: Sigma keeps moving until it finds its min Eventually $A \sim \epsilon\Lambda^4$ $m_\phi^2 = \frac{\epsilon\Lambda^4}{f^2} \sim \frac{g}{v^2}\frac{\Lambda^4}{f} = \frac{g}{f}\left(\frac{\Lambda}{v}\right)^4 v^2$

Sigma mass given by its self polynomial interaction $m_\sigma^2 \sim g_\sigma^2$

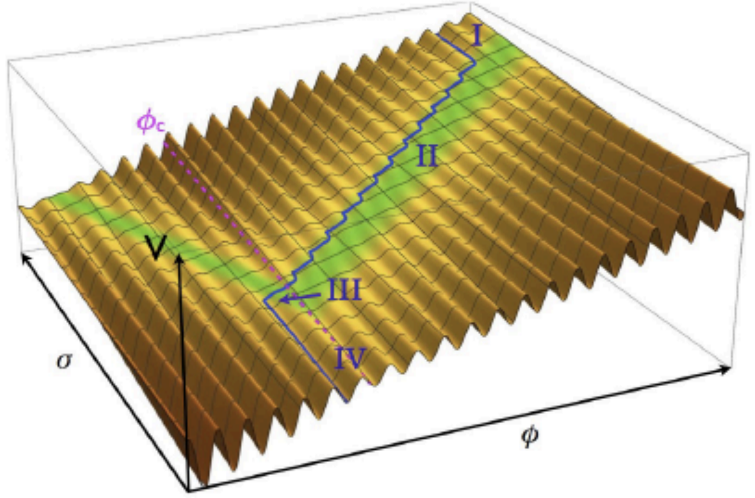
Double-scanner

$$\Delta V = g\Lambda^2\phi + g_\sigma \Lambda^2 \sigma + \left(-\Lambda^2 + g\phi\right)|h|^2 + A(\phi,\,\sigma,\,h)\cos\left(\phi/f\right)$$
$$+\frac{\partial_\mu \phi}{f_\psi}\bar{\psi}\gamma^\mu\gamma^5\psi + \frac{\partial_\mu \sigma}{f_\psi}\bar{\psi}\gamma^\mu\gamma^5\psi + m_\psi\bar{\psi}\psi\ ,$$

Constraints

$$A(\phi,\,\sigma,\,h) = \epsilon\Lambda^4\left(\beta + c_\phi\frac{g\phi}{\Lambda^2} - c_\sigma\frac{g_\sigma\sigma}{\Lambda^2} + \frac{|h|^2}{\Lambda^2}\right)$$

Periodic potential contribution to h mass-squared < v^2



$$\Delta m_h^2 \sim \epsilon\Lambda^2 \cos\left(\frac{\phi}{f}\right)_{final} \sim \epsilon\Lambda^2 \lesssim v^2$$

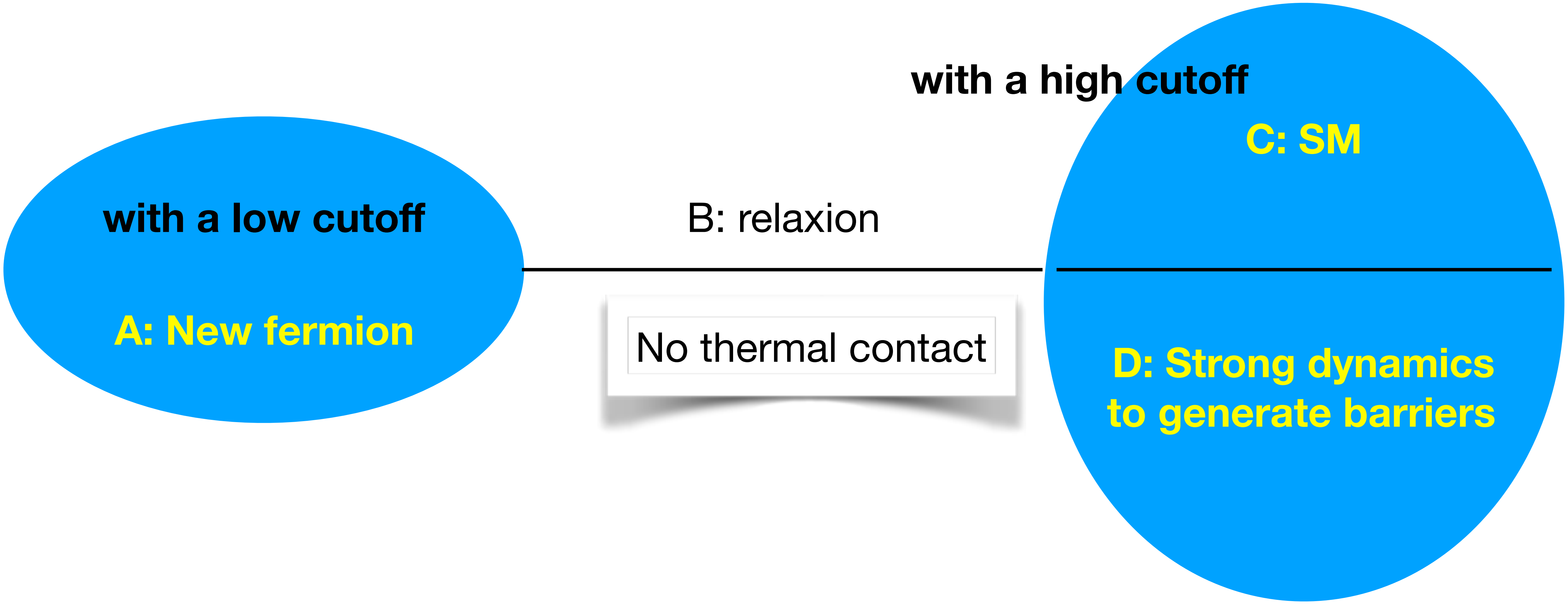
$$\Rightarrow g \lesssim \frac{v^4}{f\Lambda^2}$$

Dangerous corrections to the potential must be small

$\epsilon^2\Lambda^4\cos^2(\phi/f)$ => $\epsilon \lesssim v^2/\Lambda^2$

Further analysis and questions

$f_\psi \ll \Lambda$ universal in strong-fermion-production-supported slow-roll models



with a high cutoff

C: SM

with a low cutoff

B: relaxion

A: New fermion

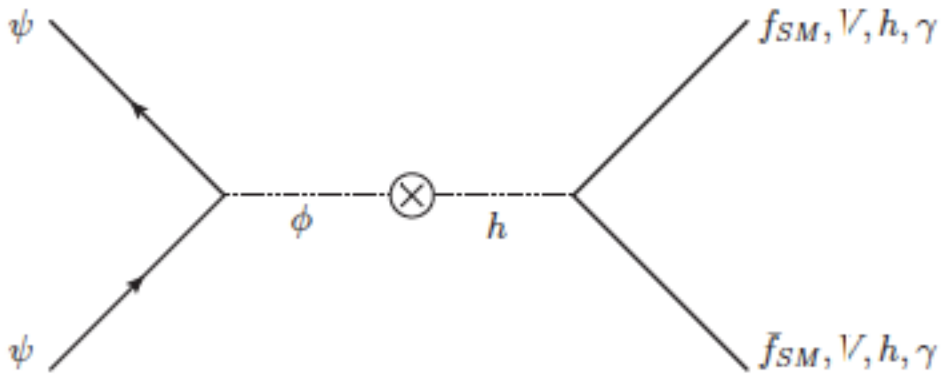
No thermal contact

D: Strong dynamics to generate barriers

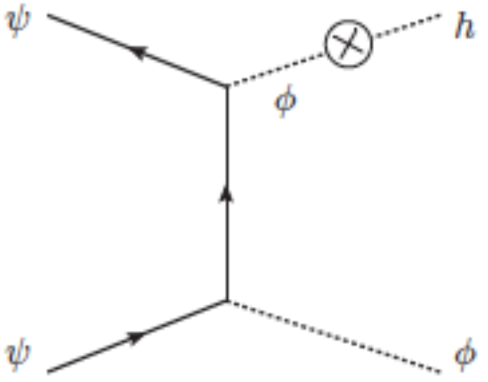
Further analysis and questions

- Can **thermal disconnection** really be true in double scanner w/o new mech.?
- An interesting question:

Chain processes: **double suppression** -> rate enough (<H)



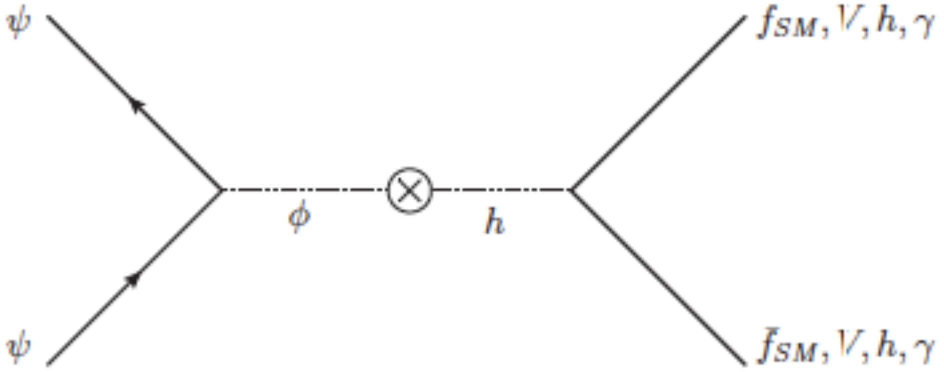
$$\Gamma \propto \mathcal{O}(g^2 v^2) \cdot \mathcal{O}(\frac{m_\psi^2}{f_\psi^2})$$



$$\Gamma \propto \mathcal{O}(g^2 v^2) \cdot \mathcal{O}(\frac{m_\psi^2}{f_\psi^4})$$

Further analysis and questions

An example of NDA analysis for a chain process



If this process is in thermal eq. when T > v (i.e. all particles are relativistic)

$$\Gamma \sim g^2 v^2 \frac{m_\psi^2}{f_\psi^2} T^{-3} > H \sim \frac{T^2}{M_P}$$

=>

$$T < \left(g^2 v^2 \frac{m_\psi^2}{f_\psi^2} M_P \right)^{1/5}$$

$$\sim 1\mathrm{GeV} \left(\frac{g}{10^{-6}\mathrm{GeV}} \right)^{2/5} \left(\frac{v}{10^{2}\mathrm{GeV}} \right)^{2/5} \left(\frac{m_\psi}{10^{-6}\mathrm{GeV}} \right)^{2/5} \left(\frac{1\mathrm{GeV}}{f_\psi} \right)^{2/5} \left(\frac{M_P}{10^{18}\mathrm{GeV}} \right)^{1/5}$$

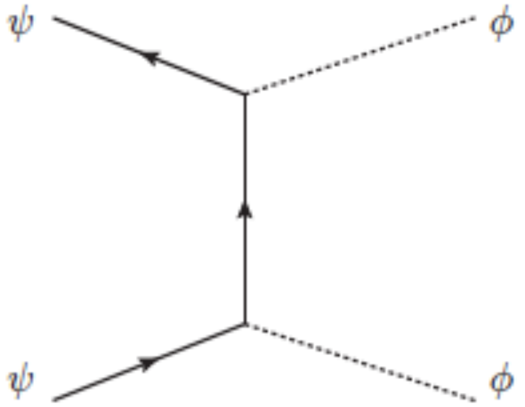
↓

contradicting w/ T > v → This chain process is not in thermal eq.

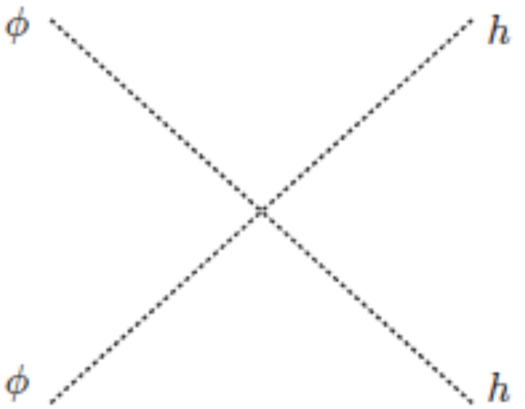
Assume:
inflaton energy dilutes faster than radiation after inflation
s.t. now the universe is radiation-dominated

Further analysis and questions

2-step processes: **single suppression** in each step, allowing each rate > H



$$\Gamma \propto \mathcal{O}\left(\frac{m_\psi^2}{f_\psi^4}\right)$$



$\epsilon h^2 \Lambda^2 \cos\left(\frac{\phi}{f}\right)$

$$\Gamma \propto \mathcal{O}\left(\epsilon^2 \frac{\Lambda^4}{f^4}\right)$$

2-step process seems to be able to make the hidden fermion and SM in eq.

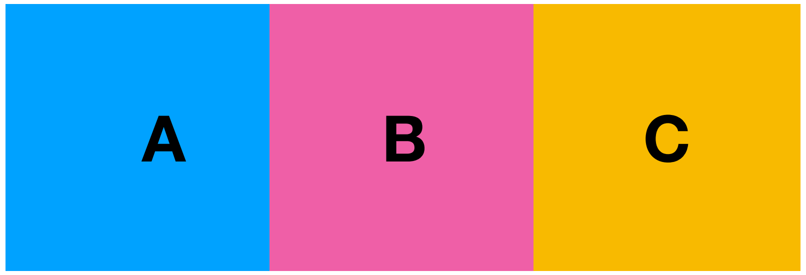
Chain VS 2-step: seemingly inconsistency

Further analysis and questions

A Chain Process VS N Processes

Further analysis and questions

- The 0th law in thermal dynamics



A <-> B, B <-> C
=>
A <->C

- But the pre-condition is, the **intermediate state** is “real”, i.e. **stable enough**

- **If B can only be off-shell, or short-lived (compared to its interaction timescales with A/C), the 2-step analysis is incorrect**

Further analysis and questions

- In our current double scanner model, T > O(100) GeV, the relaxion can be produced “on-shell” by the new fermion, and its lifetime is long enough compared to the timescales for the interactions with fermion and w/ Higgs. 2-step analysis is correct. No thermal disconnection :(

- However, the idea of “thermal-disconnection” may be applied to other models :)

- Similar problems exist in e.g. Higgs portal models. But in those models only T < (mass of intermediate particle) is interested s.t. the 2-step analysis is invalid.

Effective temperature of decoupled particles

- The momentum space distribution function after freezing-out

$$f(\vec{p},\,t) = \left[\exp\left(\frac{E-\mu}{T} \pm 1\right)\right]^{-1}$$

$$f \sim \frac{d^3 n}{dp^3}$$

$$n \sim a^{-3}$$

$$|\vec{p}| \sim a^{-1}$$

$$f \sim a^0$$

- A particle species decoupled while highly relativistic

$$E \sim |\vec{p}| \sim a^{-1} \qquad \mu \sim 0$$

$$T \sim a^{-1} \qquad T_{eff} \sim T_d \left(\frac{a_d}{a}\right)$$

- A particle species decoupled while highly non-relativistic

$$E \sim \vec{p}^2/2m \sim a^{-2} \qquad T \sim a^{-2} \qquad T_{eff} \sim T_d \left(\frac{a_d}{a}\right)^2 \qquad \mu_{eff} = m + (\mu_d - m)\frac{T_{eff}}{T_d}$$

	Pros	Cons
GKR	Start w/ EW-sym phase	Extremely small parameter, extremely large number of efolds, super-Planckian field excursion
HMT	No extremely small parameter, moderate number of efolds, sub-Planckian field excursion	specific UV needed
Ours	No extremely small parameter, moderate number of efolds, sub-Planckian field excursion, start w/ EW-sym phase	EFT consistency issue (new mech needed)

Introduction: relaxation

Problems

- Tiny coupling: e.g. g~ 10^-31 GeV for QCD relaxion
- => severe fine-tuning, exponentially large number of efolds, super-Planckian field excursion

$$\Delta\phi \geq \Lambda^2/g^2$$

↓

Contradicting with some gravity argument

Giddings and Strominger

- A free periodic scalar w/ period f has gravitational instantons S ~ M_P/f

non-negligible NP effects if f >= M_P

Whether this applies to interacting scalars: open question

Monodromy induced potential

$F_4 = dC_3$ in 4-dimensional spacetime not dynamic

$$\mathcal{L} = -\frac{1}{2}(da)^2 - V_{KS}(a) - V_{NP}(a),$$

$$V_{KS}(a) \equiv \frac{1}{2}F_4 \wedge \star_4 F_4 - mF_4 a \Rightarrow V_{KS}(a) = \frac{1}{2}(f_0 + ma)^2.$$

$$\star_4 F_4 = f_0 + ma,$$

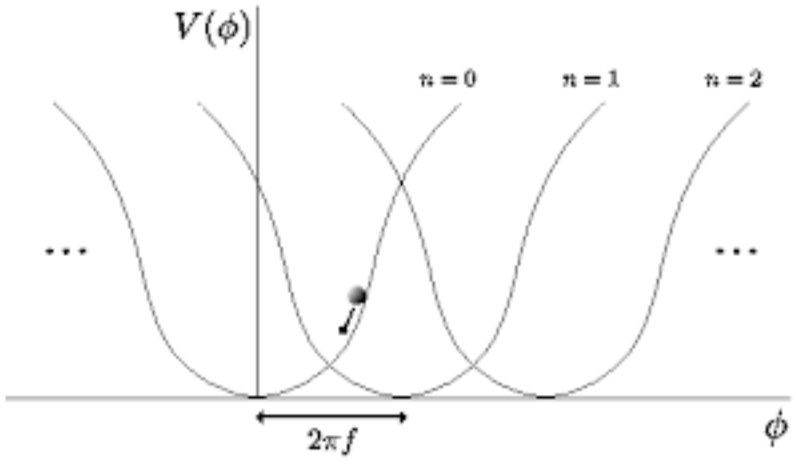
Dirac quantization of a gauge field $f_0 = n\Lambda_k^2, \quad n \in Z,$

where Λ_k is of mass dimension and the index k is associated with a combined discrete shift symmetry of the lagrangian:

$$a \to a + 2\pi f, \quad f_0 \to f_0 - 2\pi m f.$$

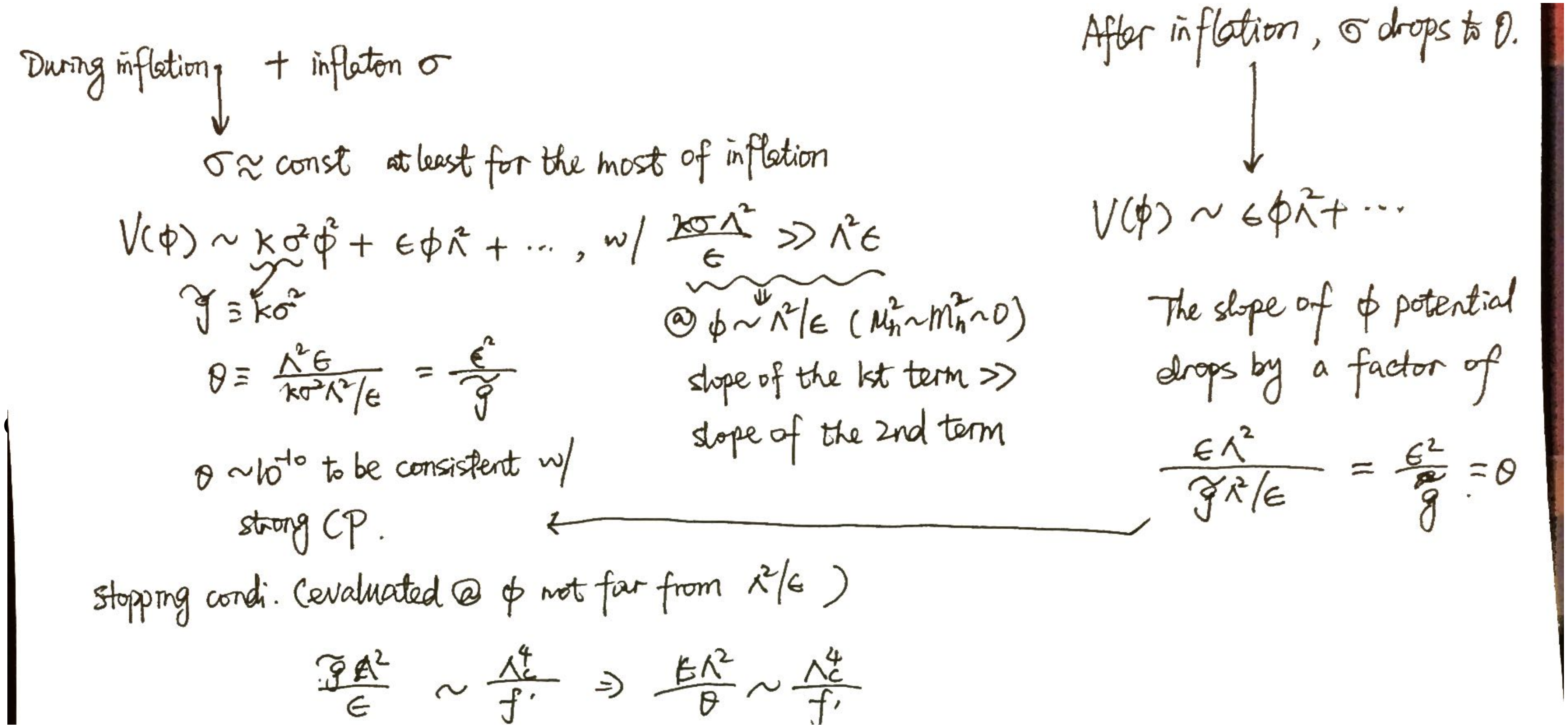
consistency condi. $2\pi m f = k\Lambda_k^2, \quad k \in Z.$

Thus the axion potential $V_{KS}(a)$ is multi-branched, with each branch (namely, a membrane) labelled by a value of f_0. When crossing a membrane, f_0 shifts by an integer times the charge of the membrane. Therefore, starting from a specific branch, the axion can go up in the potential away from its minimum and travel a distance Δa in its field space greater than the intrinsic periodicity f.



GKR’s relaxion models

- Sol: e.g. + separate inflaton, or consider non-QCD relaxion



During inflation + inflaton σ

$\sigma \approx$ const at least for the most of inflation

$V(\phi) \sim \underbrace{k\sigma^2}_{\tilde{g} \equiv k\sigma^2}\phi^2 + \epsilon\phi\Lambda^2 + \cdots$, w/ $\underbrace{\frac{k\sigma\Lambda^2}{\epsilon} \gg \Lambda^2\epsilon}$

$\theta \equiv \frac{\Lambda^2\epsilon}{k\sigma^2\Lambda^2/\epsilon} = \frac{\epsilon^2}{\tilde{g}}$

$\theta \sim 10^{-10}$ to be consistent w/ strong CP.

@ $\phi \sim \Lambda^2/\epsilon$ ($M_h^2 \sim m_h^2 \sim 0$)

slope of the 1st term $\gg$ slope of the 2nd term

After inflation, σ drops to 0.

$V(\phi) \sim \epsilon\phi\Lambda^2 + \cdots$

The slope of ϕ potential drops by a factor of

$\frac{\epsilon\Lambda^2}{\tilde{g}\Lambda^2/\epsilon} = \frac{\epsilon^2}{\tilde{g}} = \theta$

stopping condi. (evaluated @ ϕ not far from Λ^2/ϵ)

$\frac{\tilde{g}\Lambda^2}{\epsilon} \sim \frac{\Lambda_c^4}{f'} \Rightarrow \frac{\epsilon\Lambda^2}{\theta} \sim \frac{\Lambda_c^4}{f'}$

Thermal WDM

Significant difference b/t our fermion WDM and standard neutrino WDM

Ours: thermal WDM, stable

Neutrino WDM: can decay

e.g.

$N \to \nu + \gamma$ maybe detectable from X-ray observations